We compiled a sample of 200+ wide common proper motion WD+dM pairs from SLOWPoKES catalog (Dhital et al. 2010) and Silvestri et al. (2005).

**Steps to calculating WD Ages**

1. **WD atmospheric models**: Koester, private communication (Pictured above)
2. **$T_{\text{cool}}$, log g**: WD cooling tracks - Salaris et al. (2000)
3. **$M_{\text{total}}$, $t_{\text{cool}}$**: WD initial-to-final mass relationship - Catalan et al. (2008a)
4. **$M_{\text{progenitor}}$, $t_{\text{progenitor}}$**: Stellar evolution tracks - Dominguez et al. (1999)
5. **$t_{\text{progenitor}} + t_{\text{cool}}$ = Total WD age (20-50% precision)

Morgan et al. (2012) showed that WD+dM pairs with separations < 100 AU have enhanced activity above the field population - likely due to tidal effects or disk disruption. Our wide sample has separations > 1000 AU, a regime in which both components have likely evolved independently.

In our preliminary analysis we show a decreasing trend in activity strength as a function of age for dM spectral types M2-M4. Future spectroscopic observations will add to the statistical significance of this initial trend.

**Preliminary age-activity relationship**

Where SDSS spectroscopic observations in one or both components were unavailable, we performed follow-up spectroscopy at the Perkins 1.8-m telescope of Lowell Observatory. At present, 54 pairs have full spectroscopic coverage.